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V2468 CYGNI

M. L. Sitko, Space Science Institute; and D. K. Lynch, The Aerospace Corporation, report on 0.8- to 2.4-μm spectroscopy of V2468 Cyg (cf. *IAUCs* 8936, 8998) on June 18 UT using SpeX at the Infrared Telescope Facility. The object continues to develop in its coronal phase, showing lines of [Si VIII] 0.9913-μm, [Fe VI] 1.2330-μm, [Fe VI] 1.2674-μm, [Si VI] 1.964-μm, [Al IX] 2.0444-μm, [Ca VIII] 2.32-μm, and possibly [Ti VI] 1.7155-μm and [P VIII] 1.7356-μm. Many He II lines are also present. Rydberg lines at 0.8926, 1.1114, 1.5545, and 2.103 μm are also present (Lynch *et al.* 2008, *Bull. AAS* **40**, 531). These lines generally have FWHM $\sim$ 2000 km/s, and there is no indication of dust formation.

COMETS C/2009 B6, 2009 B7, 2009 C1, AND 2009 C2 (STEREO)

Further to *IAUCs* 9053, additional small and slightly diffuse Kreutz sungrazers have been found on STEREO/SECCHI HI-1A images. Very approximate peak magnitudes from K. Battams: C/2009 B6 and C/2009 B7, 11; C/2009 C1, 11-12; C/2009 C2, 10.

Comet	2009	UT	α_{2000}	δ_{2000}	Inst.	F	*MPEC*
C/2009 B6	Jan.	19.683	$23^{\rm h}21\overset{\rm m}{.}2$	$-\ 8^\circ 10'$	HI*	RM	2009-M37
C/2009 B7		29.739	23 57.1	− 3 43	HI*	RM	2009-M37
C/2009 C1	Feb.	10.044	0 37.9	+ 0 46	HI*	RK	2009-M37
C/2009 C2		11.822	0 51.3	+ 1 05	HI*	KB	2009-M37

(3200) PHAETHON

K. Battams, Naval Research Laboratory, also writes that A. Watson (Werribee, Victoria, Australia) has commented that the minor planet (3200) was visible in SECCHI HI-1A images during June 17–22, noting a very short radial elongation (perpendicular to the direction of motion) that was possibly a line-of-sight effect related to its passage through a reasonably dense, higher-speed solar outflow stream. Battams adds that the apparent brightness of (3200) increased significantly ($\sim$ 2 mag or more), peaking at mag perhaps 10–11 a few hours after perihelion (T = June 20.302 TT, q = 0.140 AU); 36 hr later, the object's had faded to magnitude roughly 13–14. Phaethon was also visible in HI-1B images during June 21–22. More formal photometry will be performed later.

