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## *SUPERNOVAE 2005kd, 2005ke, 2005kf*

Three apparent supernovae have been reported: 2005kd by T. Puckett and A. Pelloni (via the 0.35-m automated supernova patrol telescope; cf. *IAUC* 8629), and 2005ke and 2005kf by M. Baek, R. R. Prasad, and W. Li (via LOSS/KAIT; cf. *IAUC* 8628).

| SN | 2005 UT | $\alpha_{2000}$ | $\delta_{2000}$ | Mag. | *Offset* |
|---|---|---|---|---|---|
| 2005kd | Nov. 12.22 | $4^h03^m16^s.88$ | $+71^\circ43'18''.9$ | 17.0 | $0''.1$ W, $5''.0$ N |
| 2005ke | Nov. 13.33 | 3 35 04.35 | –24 56 38.8 | 17.2 | $40''$ E, $40''$ S |
| 2005kf | Nov. 11.53 | 7 47 26.51 | +26 55 32.4 | 17.2 | $0''.8$ E, $0''.6$ S |

Additional approximate magnitudes for 2005kd in PGC 14370: Sept. 18 and Nov. 9, [20.0; Nov. 13.02 UT, 16.7. Additional approximate magnitudes for 2005ke in NGC 1371: Oct. 24.32, [19.5; Nov. 14.35, 16.3. Additional approximate magnitudes for 2005kf, which is in a galaxy located $82''$ east and $16''.5$ south of the center of NGC 2449: Apr. 5.20, [19.5; Oct. 14.52, 17.0; 20.53, 16.5; 31.50, 16.7; Nov. 16.52, 17.5.

SN 2005kd has been found by an Ohio State University group (J. Eastman *et al.*, communicated by J. Prieto; cf. *CBET* 290) to be a young type-IIn supernova from a spectrogram (range 390–730 nm) taken on Nov. 13.3 UT with the MDM 2.4-m telescope (+ CCDS); the spectrum shows a blue continuum and strong hydrogen Balmer (H$\alpha$, H$\beta$, H$\gamma$, and H$\delta$) and He I lines in emission at central wavelengths fully consistent with the recession velocity of PGC 14370.

## *(2006) POLONSKAYA*

D. Pray, Greene, RI; P. Pravec and P. Kušnirák, Ondřejov Observatory; and W. Cooney, J. Gross and D. Terrell, Sonoita Research Observatory, report that observations obtained during Nov. 1–13 show that the minor planet (2006) consists of a pair of bodies orbiting each other with a period of 19.15 $\pm$ 0.02 hr and producing mutual eclipse/occultation events 0.06-mag deep. There are present two rotational lightcurves with periods of 3.1183 $\pm$ 0.0002 and 6.656 $\pm$ 0.001 hr, with amplitudes of 0.08 and 0.06 mag, respectively, that persist during the mutual events. This suggests that either the system is binary with both bodies rotating non-synchronously with the orbital motion or there is a third body producing the additional period. A lower limit on the mean-diameter ratio of the eclipsing pair is 0.22. Further observations are needed to fully resolve the system.

2005 November 16  *Daniel W. E. Green*